Blazar Heating – The Rosetta Stone for Structure Formation?

Christoph Pfrommer[1]

in collaboration with

Avery E. Broderick[2], Phil Chang[2], Ewald Puchwein[1], Volker Springel[1]

[1] Heidelberg Institute for Theoretical Studies, Germany
[2] Canadian Institute for Theoretical Astrophysics, Canada

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Outline

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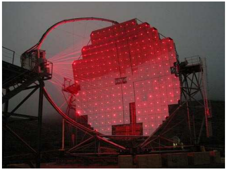
1. Physics of blazar heating
   - TeV emission from blazars
   - Propagation of TeV photons
   - Plasma instabilities
2. The intergalactic medium
   - Properties of blazar heating
   - Thermal history of the IGM
   - The Lyman-α forest
3. Structure formation
   - Entropy evolution
   - Bimodality of galaxy clusters
   - Formation of dwarf galaxies

TeV gamma-ray astronomy

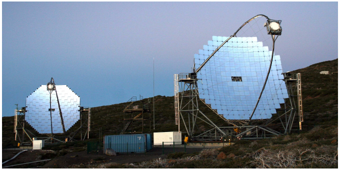
H.E.S.S.

MAGIC I



VERITAS

MAGIC II

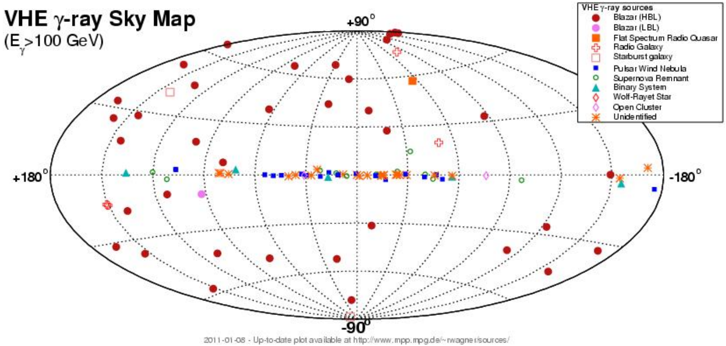


HITS

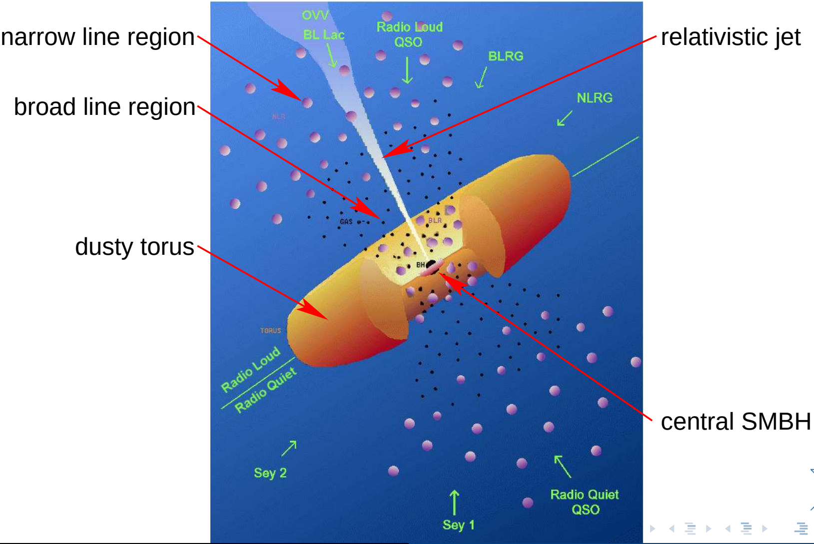
The TeV gamma-ray sky

There are several classes of TeV sources:

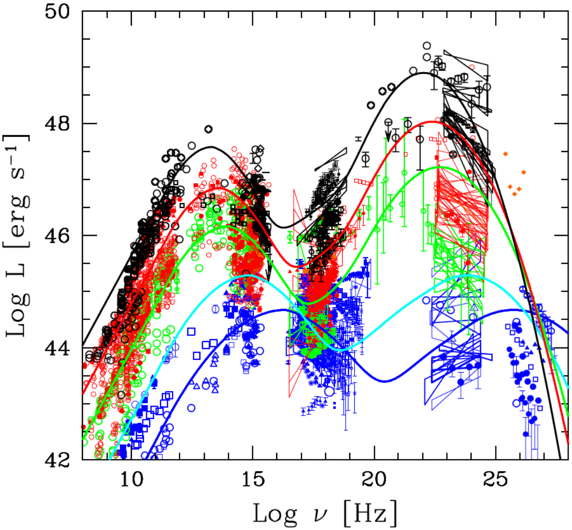
- Galactic - pulsars, BH binaries, supernova remnants
- Extragalactic - **mostly** blazars, two starburst galaxies



Unified model of active galactic nuclei



The blazar sequence



Ghisellini (2011)

Propagation of TeV photons

- 1 TeV photons can pair produce with 1 eV photons:

$$\gamma + \gamma \rightarrow e^{+} + e^{-}$$

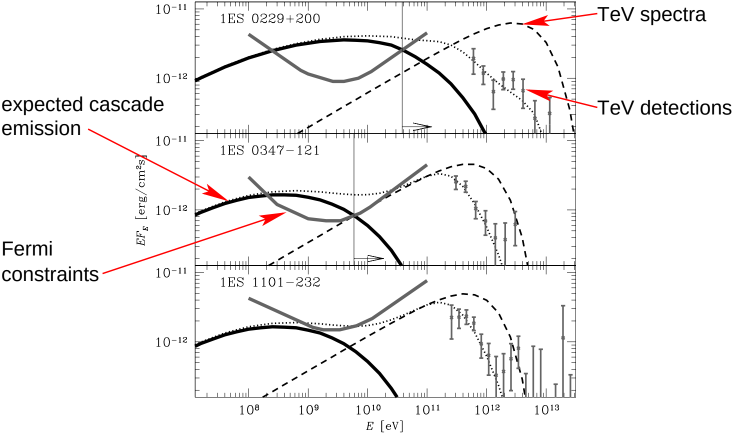
- mean free path for this depends on the density of 1 eV photons:
  $\rightarrow$ typically $\sim$ 100 Mpc
  $\rightarrow$ pairs produced with energy of 0.5 TeV ($\gamma = 10^6$)
- these pairs inverse Compton scatter off the CMB photons
  $\rightarrow$ mean free path is $\sim$ 30 kpc
  $\rightarrow$ producing gamma-rays of $\sim$ 1 GeV

$$E \sim \gamma^2 E_{\mathrm{CMB}} \sim 1\ \mathrm{GeV}$$

- each TeV point source is also a GeV point source

What about the cascade emission?

Every TeV source should be associated with a 1-100 GeV gamma-ray halo – **not seen!**

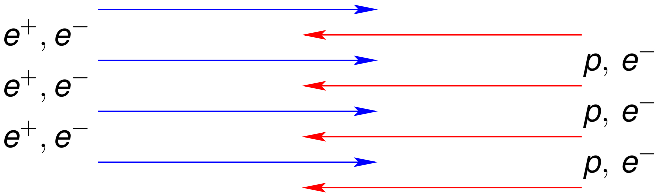


Neronov & Vovk (2010)

Missing plasma physics?

How do beams of e^+/e^- propagate through the IGM?

- plasma processes are important
- interpenetrating beams of charged particles are unstable
- consider the two-stream instability for two beams:



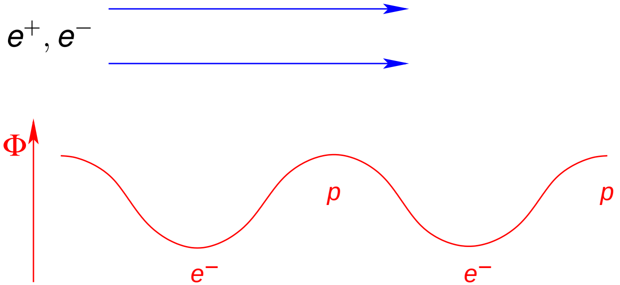
- one frequency (timescale) and one length in the problem:

$$\frac{\omega_p}{\gamma} = \sqrt{\frac{4\pi e^2 n_e}{\gamma^2 m_e}} \qquad \lambda_p = \frac{\gamma c}{\omega_p}$$

Two-stream instability: mechanism

wave-like perturbation with $\boldsymbol{k} || \boldsymbol{v}_{\text{beam}}$, longitudinal charge oscillations in background plasma (Langmuir wave):

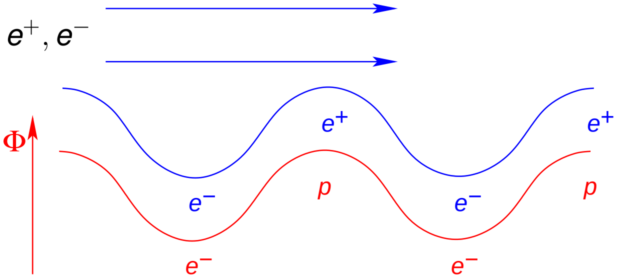
- initially homogeneous beam-e^-:
  attractive (repulsive) force by potential maxima (minima)
- e^- attain lowest velocity in potential minima → bunching up
- e^+ attain lowest velocity in potential maxima → bunching up



Two-stream instability: mechanism

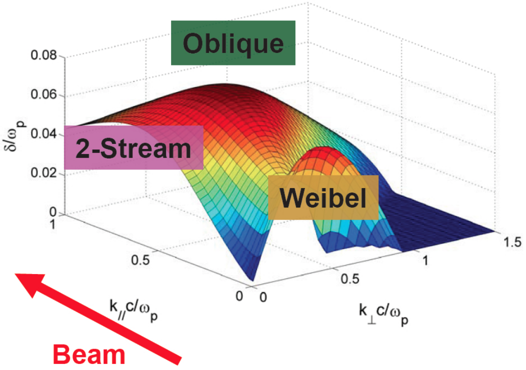
wave-like perturbation with $\boldsymbol{k} || \boldsymbol{v}_{\rm beam}$, longitudinal charge oscillations in background plasma (Langmuir wave):

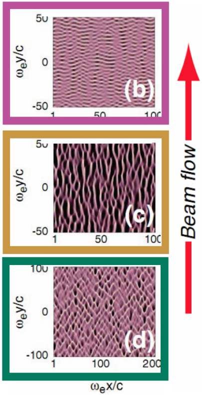
- beam-e^+/e^- couple in phase with the background perturbation: enhances background potential
- stronger forces on beam-e^+/e^- $\rightarrow$ positive feedback
- exponential wave-growth $\rightarrow$ instability



Oblique instability

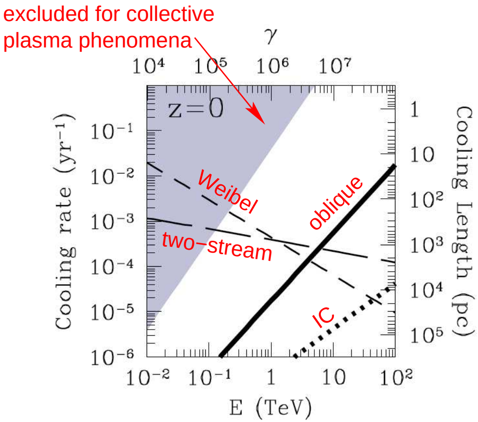
k oblique to $\boldsymbol{v}_{\rm beam}$: real word perturbations don't choose "easy" alignment $= \sum$ all orientations





Bret (2009), Bret+ (2010)

Beam physics – growth rates



Broderick, Chang, C.P. (2011)

- consider a light beam penetrating into relatively dense plasma
- maximum growth rate

$$\sim 0.4\,\gamma\,\frac{n_{\rm beam}}{n_{\rm IGM}}\,\omega_p$$

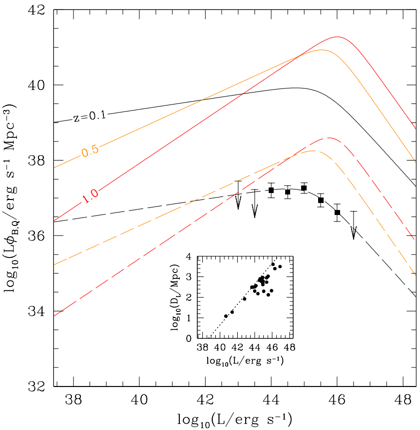
- oblique instability beats IC by two orders of magnitude

Beam physics – growth rates

- non-linear evolution of these instabilities at these density contrasts is not known
- expectation from PIC simulations suggest substantial isotropization of the beam
- plasma instabilities cool the beam, no energy left over for IC off the CMB

Outline

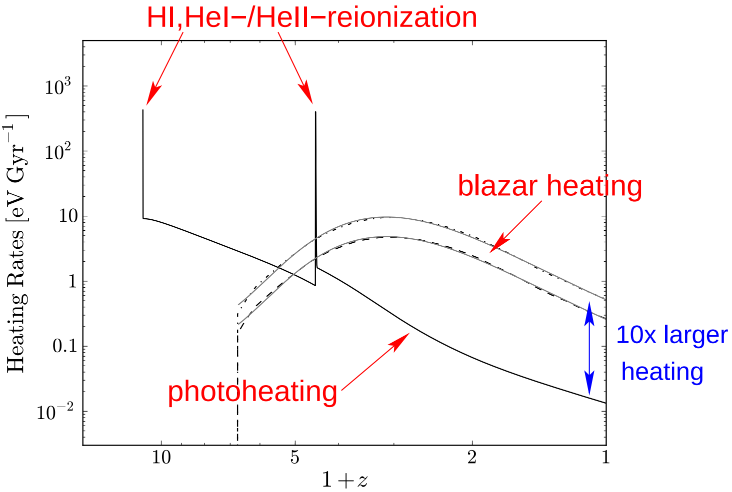
TeV blazar luminosity density



Broderick, Chang, C.P. (2011)

- collect luminosity of all 23 TeV blazars with good spectral measurements
- account for the selection effects
- TeV blazar luminosity density is a scaled version ($\sim$ 0.2%) of that of quasars!
- assume that they trace each other for all *z*

Evolution of the heating rates



Chang, Broderick, C.P. (2011)

Blazar heating vs. photoheating

- total power from AGN/stars vastly exceeds the TeV power of blazars
- $T_{\mathrm{IGM}} \sim 10^4$ K (1 eV) at mean density ($z \sim 2$)

$$\varepsilon_{\mathrm{th}} = \frac{kT}{m_{\mathrm{p}} c^2} \sim 10^{-9}$$

- radiative energy ratio emitted by BHs in the Universe (Fukugita & Peebles 2004)

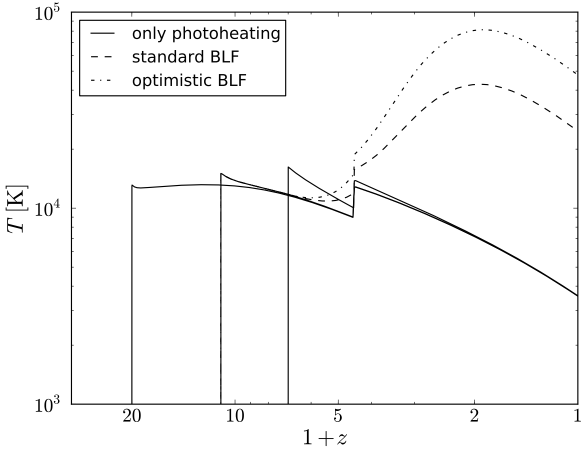
$$\varepsilon_{\mathrm{rad}} = \eta\, \Omega_{\mathrm{bh}} \sim 0.1 \times 10^{-4} \sim 10^{-5}$$

- fraction of the energy energetic enough to ionize H I is ~ 0.1:

$$\varepsilon_{\mathrm{UV}} \sim 0.1 \varepsilon_{\mathrm{rad}} \sim 10^{-6} \quad \rightarrow \quad kT \sim \mathrm{keV}$$

- photoheating efficiency $\eta_{\mathrm{ph}} \sim 10^{-3} \quad \rightarrow \quad kT \sim \eta_{\mathrm{ph}}\, \varepsilon_{\mathrm{UV}}\, m_{\mathrm{p}} c^2 \sim \mathrm{eV}$
  (limitted by the abundance of H I/He II due to the small recombination rate)
- blazar heating efficiency $\eta_{\mathrm{bh}} \sim 10^{-3} \quad \rightarrow \quad kT \sim \eta_{\mathrm{bh}}\, \varepsilon_{\mathrm{rad}}\, m_{\mathrm{p}} c^2 \sim 10\,\mathrm{eV}$
  (limited by the total power of TeV sources)

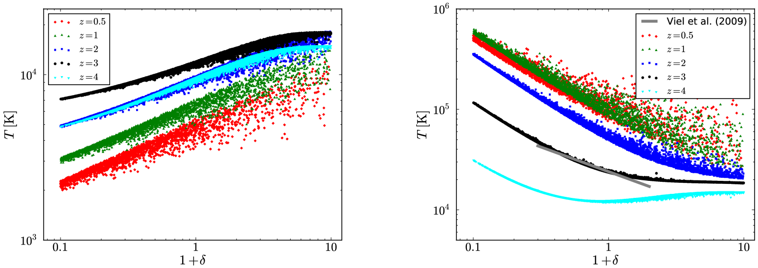
Thermal history of the IGM



Chang, Broderick, C.P. (2011)

Evolution of the equation of state

no blazar heating | blazar heating

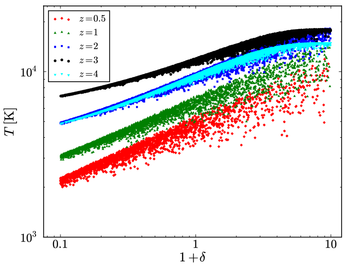


Chang, Broderick, C.P. (2011)

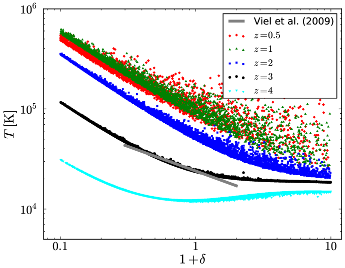
- blazars and extragalactic background light are uniform
  → blazar heating independent of density
  → causes inverted equation of state, $T \propto 1/\delta$
- blazars completely change the thermal history of the diffuse IGM and late-time structure formation

Evolution of the equation of state

no blazar heating



blazar heating



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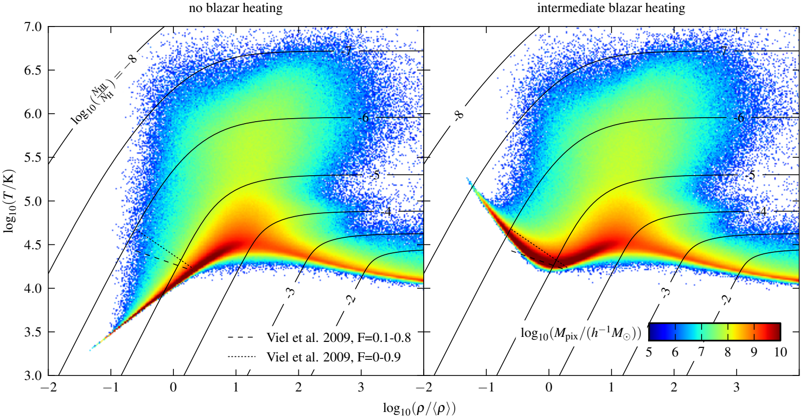
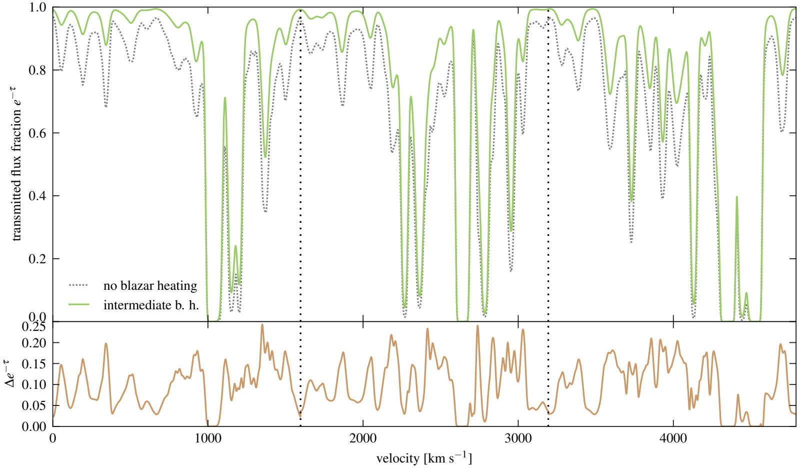
Equation of state



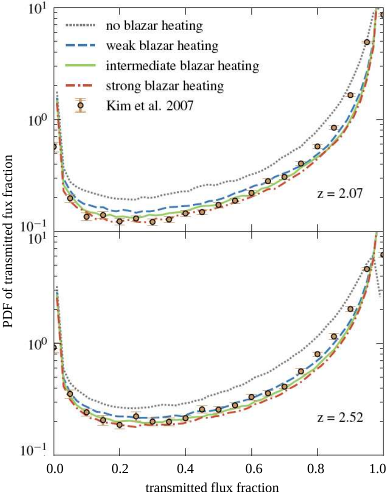
Puchwein, C.P., Springel, Broderick, Chang (2011)

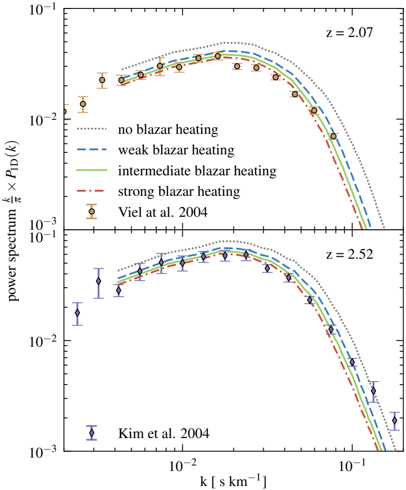
Ly-α spectra



Puchwein+ (2011)

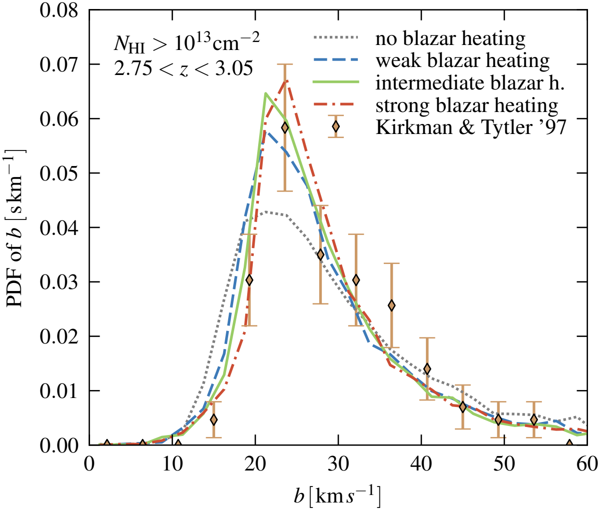
Ly-α flux PDFs and power spectra





Puchwein+ (2011)

Voigt profile fitting – line width distribution



Puchwein+ (2011)

Lyman-α forest in a blazar heated Universe

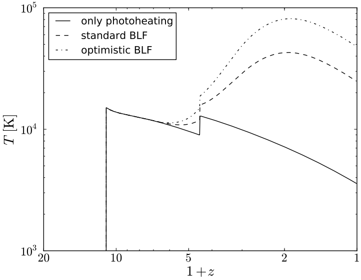
impressive improvement in modelling the Lyman-α forest is a direct consequence of the peculiar properties of blazar heating:

- heating rate independent of IGM density $\rightarrow$ naturally produces the inverted EOS that Lyman-α forest data demand
- recent and continuous nature of the heating needed to match the redshift evolutions of all Lyman-α forest statistics
- magnitude of the heating rate required by Lyman-α forest data $\sim$ the total energy output of TeV blazars (or equivalently $\sim 0.2\%$ of that of quasars)

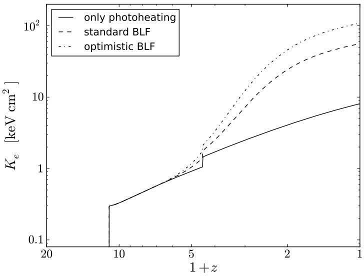
Outline

Entropy evolution

temperature evolution



entropy evolution

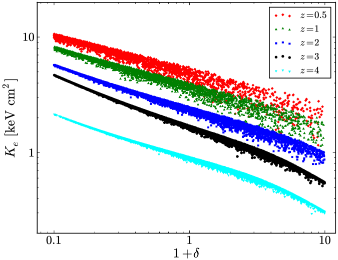


C.P., Chang, Broderick (2011)

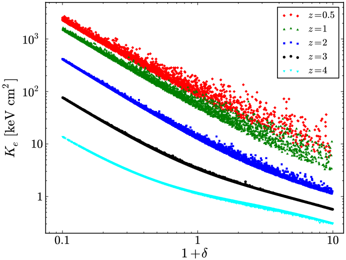
- evolution of the entropy, $K_\mathrm{e} = kTn_\mathrm{e}^{-2/3}$, at mean density
- blazar heating substantially increases the entropy floor ($z \lesssim 2$)

Evolution of the entropy equation of state

no blazar heating



blazar heating



C.P., Chang, Broderick (2011)

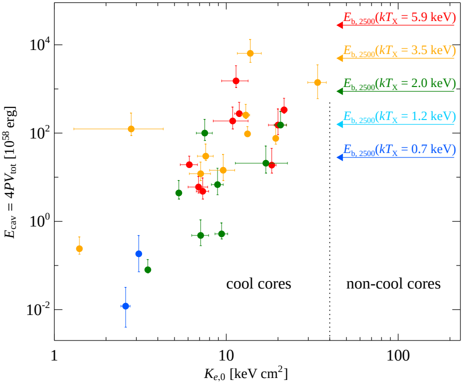
- blazar heating substantially increases the entropy in voids
- scatter is also increased → larger stochasticity of structure formation

Blazar heating: AGN feedback vs. pre-heating

Blazar heating is an amalgam of pre-heating and AGN feedback:

- blazar heating is not localized ($\neq$ AGN feedback)
  $\rightarrow$ changes initial conditions for forming groups (but provides no stability for cool cores, CCs)
- blazar heating generates time-dependent entropy floor ($\neq$ pre-heating)
  $\rightarrow$ solves the classical problems of pre-heating ($z \sim 3$):
  - provides a physical mechanism
  - does not starve galaxy formation for $z \lesssim 3$
  - early forming groups can cool and develop observed low-K_e cores

How efficient is heating by AGN feedback?



C.P., Chang, Broderick (2011)

- cavity enthalpy

$$E_{cav} = 4\,PV_{tot}$$

- in some cases

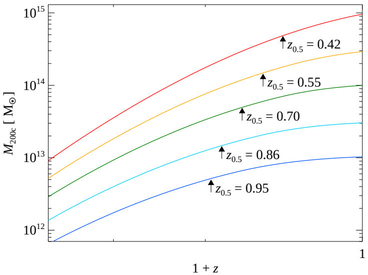
$$E_{cav} \gtrsim E_{bind}(R_{2500})$$

- cavity energy only couples weakly into ICM, but prevents cooling catastrophe

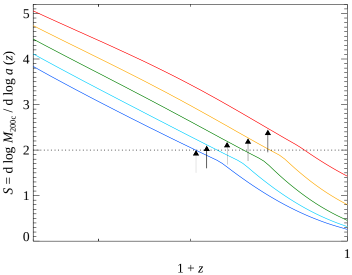
- on a buoyancy timescale, no AGN outburst transforms a CC to a non-cool core (NCC) cluster!

Mass accretion history of groups/clusters

mass accretion history



mass accretion rates

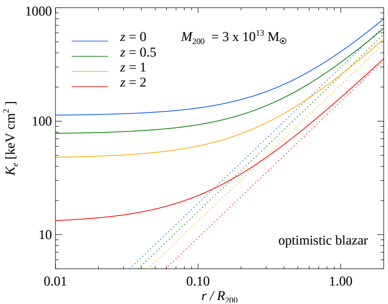


C.P., Chang, Broderick (2011)

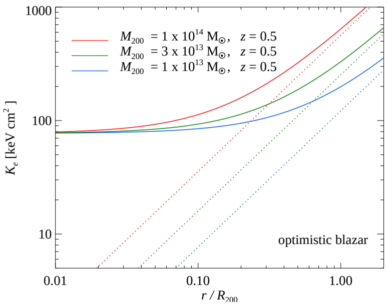
- peak entropy injection from blazar heating ($z \sim 1$) matches formation time of groups
- early forming groups are unaffected and develop cool cores
- late forming groups have an elevated entropy core

Entropy profiles: effect of blazar heating

varying formation time



varying cluster mass



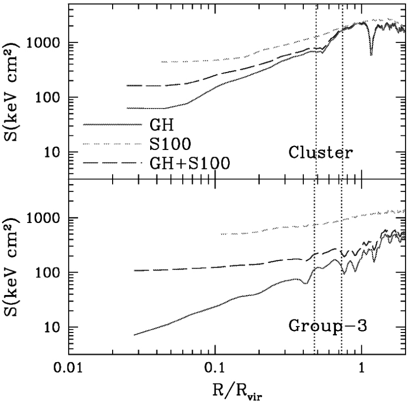
C.P., Chang, Broderick (2011)

- cluster entropy profile immediately after formation (no cooling)
- largest effect for late forming, small objects

Scenario for the bimodality of cluster core entropies ?

- entropy core, $K_{e,0}$, immediately after formation is set by the z-dependent blazar heating
- only late forming groups ($z \lesssim 1$) are directly affected by blazar (pre-)heating
- if the cooling time, t_{cool}, is shorter than the time period to the successive merger, t_{merger}, the group will radiate away the elevated core entropy and evolve into a CC
- if $t_{cool} > t_{merger}$, merger shocks can gravitationally reprocess the entropy cores and amplify them → potentially those forming clusters evolve into non-cool core (NCC) systems

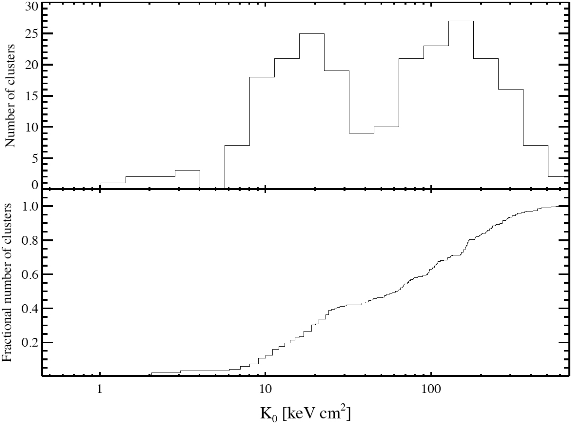
Gravitational reprocessing of entropy floors



Borgani+ (2005)

- larger $K_{e,0}$ of a merging cluster facilitates shock heating $\rightarrow$ increase of $K_{e,0}$ over entropy floor
- entropy floor of 100 keV cm^2 at $z = 3$ in non-radiative simulation:
  net entropy amplification factor $\sim$ 3–5 for clusters and groups (Borgani+ 2005)
- expect median of $K_{e,0} \sim 150$ keV cm^2; maximum $K_{e,0} \sim 600$ keV cm^2

Bimodality of cluster core entropies



Cavagnolo+ (2009)

- *Chandra* observations match blazar heating expectations!
- need hydrodynamic simulations to confirm this scenario

Jeans mass

- on small enough scales, the thermal pressure can oppose gravitational collapse of the gas
- characteristic length scale below which objects will not form
- Jeans wavenumber and mass is obtained by balancing the sound crossing and free-fall timescales

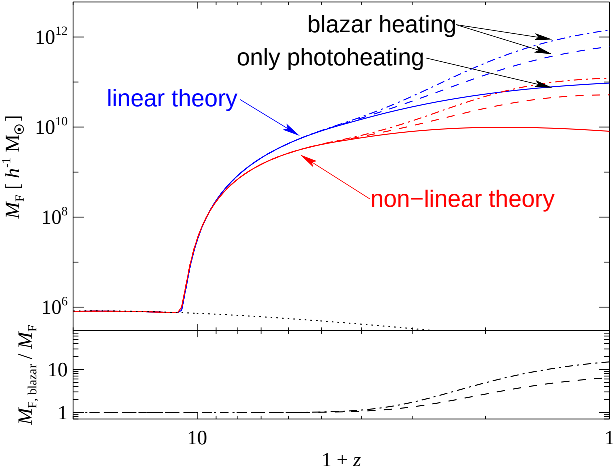
$$k_J(a) \equiv \frac{a}{c_s(a)} \sqrt{4\pi G \bar{\rho}(a)}$$

$$M_J(a) \equiv \frac{4\pi}{3} \bar{\rho}(a) \left(\frac{2\pi a}{k_J(a)}\right)^3 = \frac{4\pi^{5/2}}{3} \frac{c_s^3(a)}{G^{3/2} \bar{\rho}^{1/2}(a)}$$

- blazar heating increases the IGM temperature by ~ 10:

$$\frac{M_{J,\mathrm{blazar}}}{M_{J,\mathrm{photo}}} = \left(\frac{c_{\mathrm{s,blazar}}}{c_{\mathrm{s,photo}}}\right)^3 = \left(\frac{T_{\mathrm{blazar}}}{T_{\mathrm{photo}}}\right)^{3/2} \gtrsim 30$$

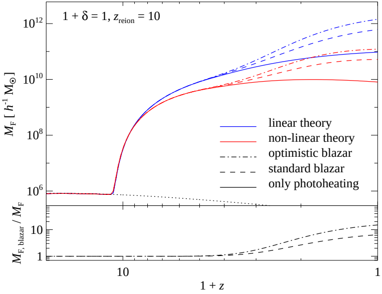
Filtering mass – dwarf formation



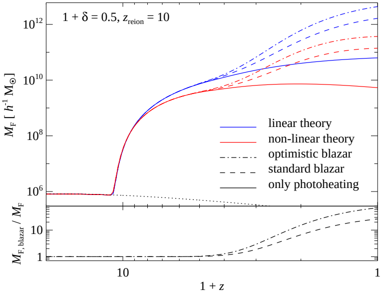
C.P., Chang, Broderick (2011)

Peebles' void phenomenon explained?

mean density



void, $1 + \delta = 0.5$

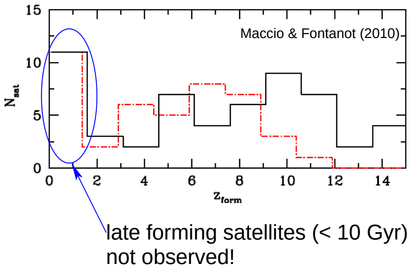


C.P., Chang, Broderick (2011)

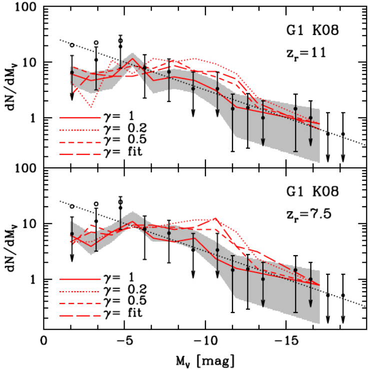
- blazar heating efficiently suppresses the formation of void dwarfs within existing DM halos of masses $< 3 \times 10^{11}$ M$_\odot$ ($z = 0$)
- reconciling the number of void dwarfs in simulations and the paucity of those in observations

"Missing satellite" problem in the Milky Way

satellite formation time



satellite luminosity function



Maccio+ (2010)

- blazar heating suppresses late satellite formation, reconciling low observed dwarf abundances with CDM simulations

Conclusions on blazar heating

- novel mechanism; dramatically alters thermal history of the IGM:
  - uniform and z-dependent preheating
  - rate independent of density $\rightarrow$ inverted EOS
  - consistent picture of Lyman-α forest
- significantly modifies late-time structure formation:
  - group/cluster bimodality of core entropy values
  - may suppress Sunyaev-Zel'dovich power spectrum
  - dwarf formation: "missing satellite" problem, void phenomenon
- explains puzzles in high-energy astrophysics:
  - TeV blazars can evolve like quasars
  - extragalactic gamma-ray background at $E \gtrsim 10$ GeV
  - invalidates intergalactic B-constraints from blazar spectra