Cosmic ray feedback in galaxy formation

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in collaboration with

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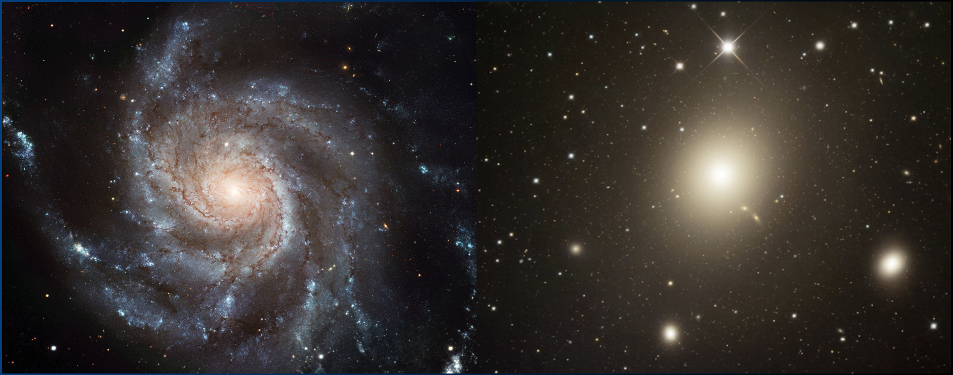
Magnetic fields in galaxies - RU Bochum - June 2016

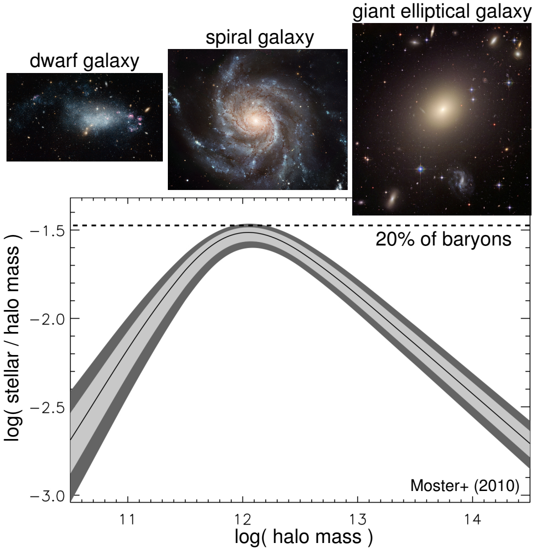
Outline

1. Introduction and Motivation
    - Puzzles in galaxy formation
    - Galactic winds
    - Cosmic rays
2. Cosmic ray simulations
    - Sedov explosions
    - Galaxy simulations
    - Cosmological simulations
3. AGN feedback
    - Radio and γ-ray emission
    - Cosmic-ray heating
    - Conclusions

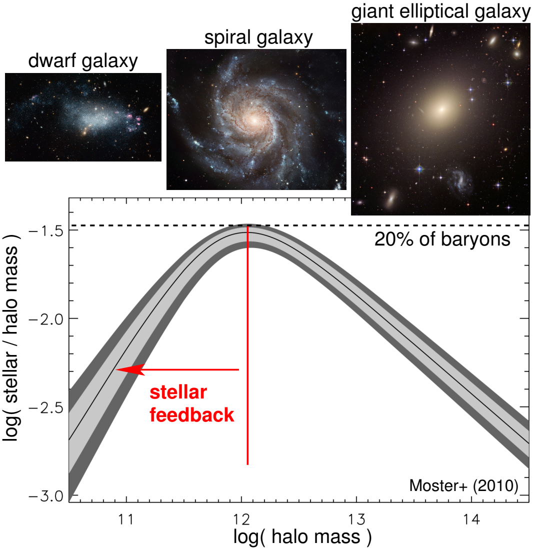
Puzzles in galaxy formation



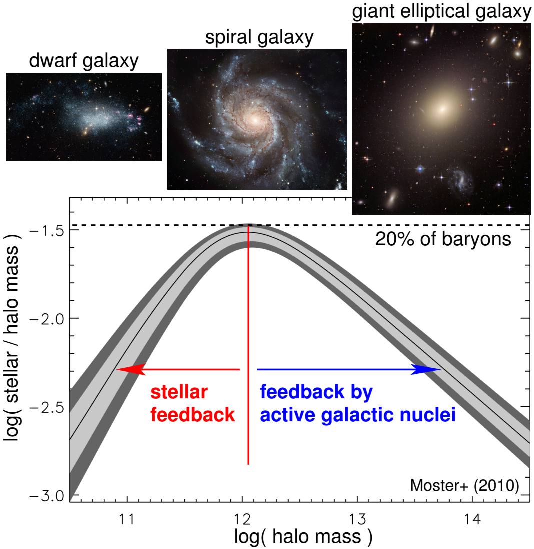
Puzzles in galaxy formation



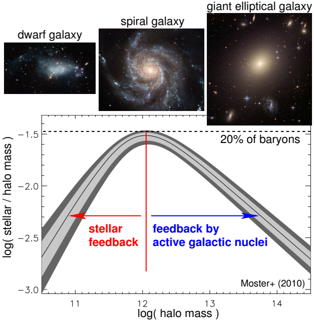
Puzzles in galaxy formation



Puzzles in galaxy formation



Puzzles in galaxy formation



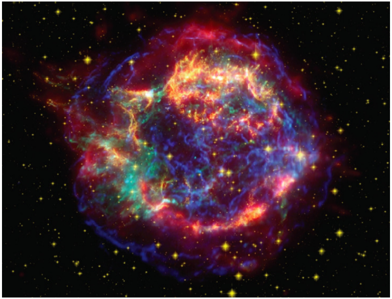
Bright-end of luminosity function:

- astrophysical solutions:
  AGN/quasar feedback, . . .

Faint-end of luminosity function:

- dark matter (DM) solutions:
  warm DM, interacting DM, DM from late decays, large annihilation rates, . . .
- astrophysical solutions:
  - preventing gas from falling into DM potential wells:
    increasing entropy by reionization, blazar heating . . .
  - preventing gas from forming stars in galaxies:
    suppress cooling (photoionization, low metallicities), . . .
  - pushing gas out of galaxies:
    supernova/quasar feedback → **galactic winds**

Galactic winds



supernova Cassiopeia A

X-ray: NASA/CXC/SAO; Optical: NASA/STScI;
Infrared: NASA/JPL-Caltech/Steward/O.Krause et al.

- galactic supernova remnants drive shock waves, turbulence, accelerate electrons + protons, amplify magnetic fields

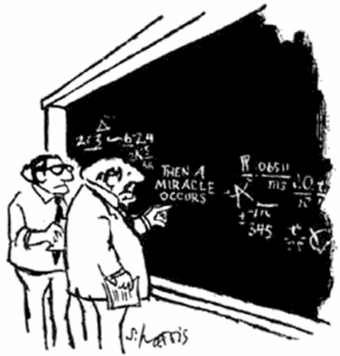
Galactic winds

super wind in M82

NASA/JPL-Caltech/STScI/CXC/UofA

- galactic supernova remnants drive shock waves, turbulence, accelerate electrons + protons, amplify magnetic fields
- star formation and supernovae drive gas out of galaxies by galactic super winds
- critical for understanding the physics of galaxy formation → may explain puzzle of low star conversion efficiency in dwarf galaxies

Galactic winds



"I THINK YOU SHOULD BE MORE EXPLICIT HERE IN STEP TWO."

© Sydney Harris

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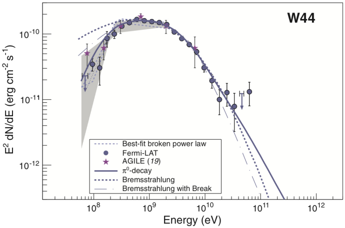
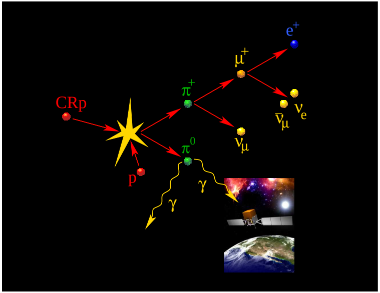
The role of supernova remnants

- supernova remnant shocks amplify magnetic fields and accelerate CR electrons up to $\sim$ 100 TeV (narrow X-ray synchrotron filaments observed by *Chandra*)
- pion bump provides evidence for CR proton acceleration (*Fermi*/AGILE γ-ray spectra)

Fermi observations of W44:



Ackermann+ (2013)



The role of supernova remnants

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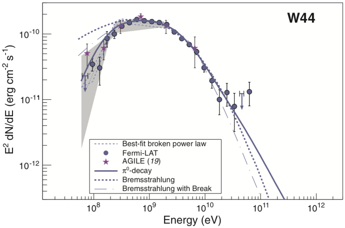
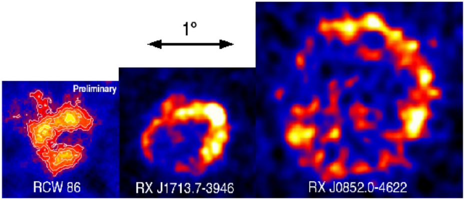
- pion bump provides evidence for CR proton acceleration (*Fermi*/AGILE γ-ray spectra)
- shell-type SNRs show evidence for efficient shock acceleration beyond $\sim$ 100 TeV (HESS TeV γ-ray observations)

Fermi observations of W44:



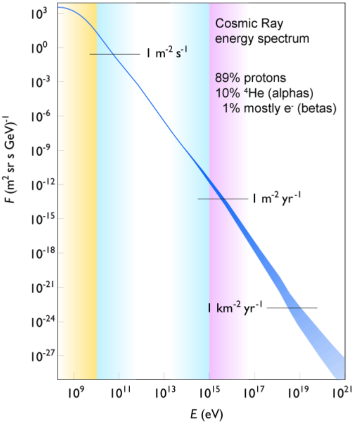
Ackermann+ (2013)

HESS observations of shell-type SNRs:



Hinton (2009)

Galactic cosmic ray spectrum

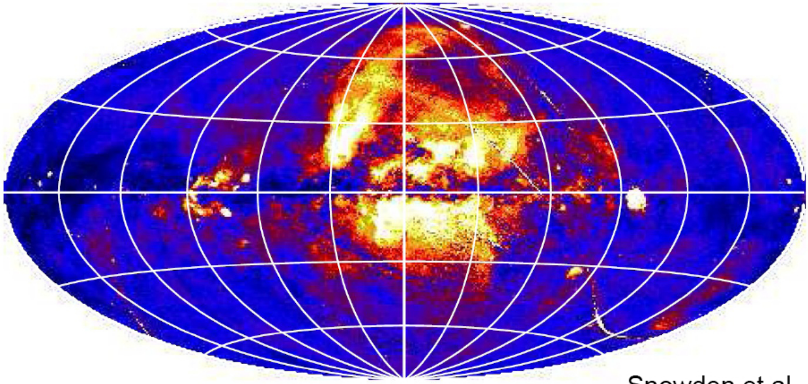


data compiled by Swordy

- spans more than 33 decades in flux and 12 decades in energy
- “knee” indicates characteristic maximum energy of galactic accelerators
- CRs beyond the “ankle” have extra-galactic origin
- energy density of cosmic rays, magnetic fields, and turbulence in the interstellar gas all similar

Galactic wind in the Milky Way?

Diffuse X-ray emission in our galaxy

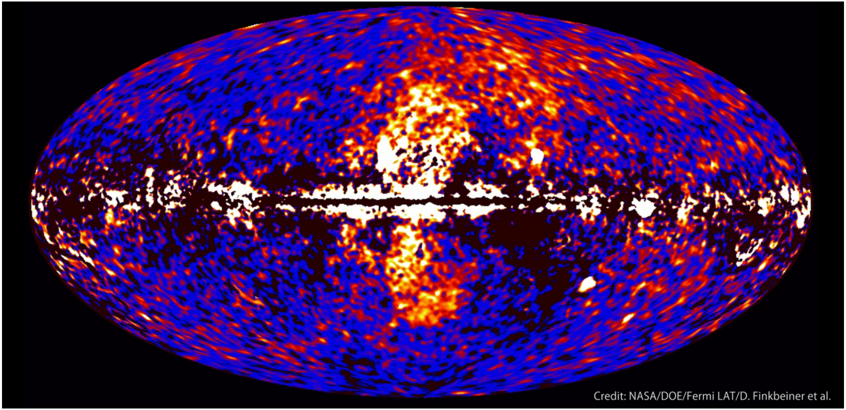


Snowden et al., 2007

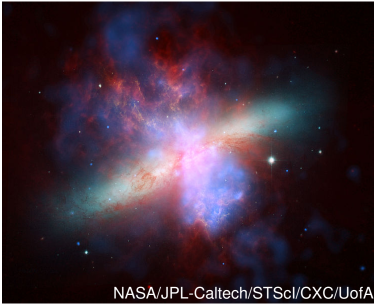
. . . as suggested by Everett+ (2008) and Everett, Schiller, Zweibel (2010)

Galactic wind in the Milky Way?

Fermi gamma-ray bubbles


Credit: NASA/DOE/Fermi LAT/D. Finkbeiner et al.

How are galactic winds driven?


NASA/JPL-Caltech/STScI/CXC/UofA

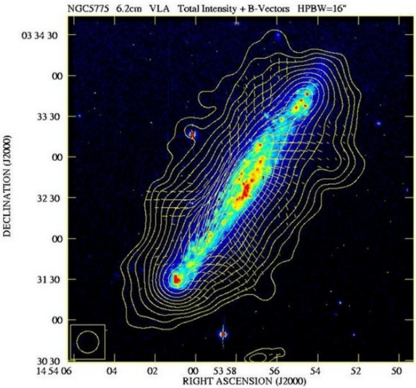
super wind in M82

- thermal pressure provided by supernovae or AGNs?
- radiation pressure and photoionization by massive stars and QSOs?
- cosmic-ray (CR) pressure and Alfvén wave heating of CRs accelerated at supernova shocks?

observed energy equipartition between cosmic rays, thermal gas and magnetic fields
→ suggests self-regulated feedback loop with CR driven winds

Why are CRs important for wind formation?

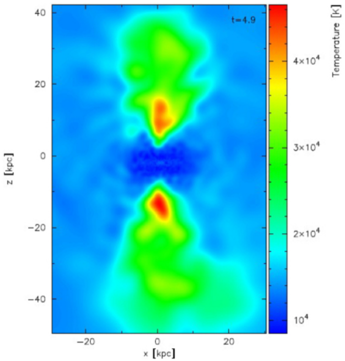
Radio halos in disks: CRs and magnetic fields exist at the disk-halo interface



Tüllmann+ (2000)

- CR pressure drops less quickly than thermal pressure ($P \propto \rho^{\gamma}$)
- CRs cool less efficiently than thermal gas
- CR pressure energizes the wind → “CR battery”
- poloidal (“open”) field lines at wind launching site → CR-driven Parker instability

Cosmic-ray driven winds – literature

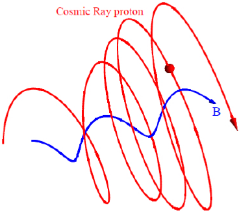


Uhlig, C.P.+ (2012)

- previous theoretical works: Ipavich (1975), Breitschwerdt+ (1991), Zirakashvili+ (1996), Ptuskin+ (1997), Breitschwerdt+ (2002), Socrates+ (2008), Everett+ (2008, 2010), Samui+ (2010), Dorfi & Breitschwerdt (2012)
- previous 3D simulations:
  CR streaming: Uhlig, C.P.+ (2012)
  CR diffusion: Booth+ (2013), Hanasz+ (2013), Salem & Bryan (2014)

Interactions of CRs and magnetic fields

- CRs scatter on magnetic fields → isotropization of CR momenta
- **CR streaming instability:** Kulsrud & Pearce 1969
  - if $v_{\rm cr} > v_{\rm A}$, CR current provides steady driving force, which amplifies an Alfvén wave field in resonance with the gyroradii of CRs
  - scattering off of this wave field limits the (GeV) CRs' bulk speed $\sim v_{\rm A}$
  - wave damping: transfer of CR energy and momentum to the thermal gas



→ **CRs exert a pressure on the thermal gas by means of scattering off of Alfvén waves**

CR transport

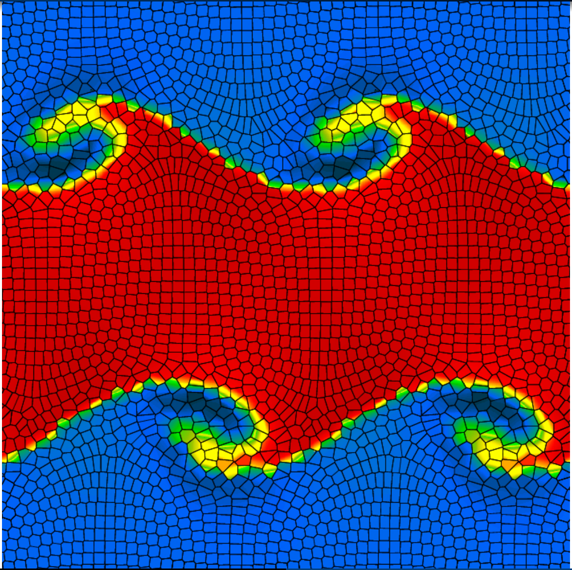
- total CR velocity $\boldsymbol{v}_{\mathrm{cr}} = \boldsymbol{v} + \boldsymbol{v}_{\mathrm{st}} + \boldsymbol{v}_{\mathrm{di}}$ (where $\boldsymbol{v} \equiv \boldsymbol{v}_{\mathrm{gas}}$)
- CRs stream down their own pressure gradient relative to the gas, CRs diffuse in the wave frame due to pitch angle scattering by MHD waves (both transports are along the local direction of $\boldsymbol{B}$):

$$\boldsymbol{v}_{\mathrm{st}} = -\frac{\boldsymbol{B}}{\sqrt{4\pi\rho}} \frac{\boldsymbol{b}\cdot\boldsymbol{\nabla} P_{\mathrm{cr}}}{|\boldsymbol{b}\cdot\boldsymbol{\nabla} P_{\mathrm{cr}}|}, \qquad \boldsymbol{v}_{\mathrm{di}} = -\kappa_{\mathrm{di}}\boldsymbol{b}\,\frac{\boldsymbol{b}\cdot\boldsymbol{\nabla}\varepsilon_{\mathrm{cr}}}{\varepsilon_{\mathrm{cr}}},$$

- energy equations with $\varepsilon = \varepsilon_{\mathrm{th}} + \rho v^2/2$:

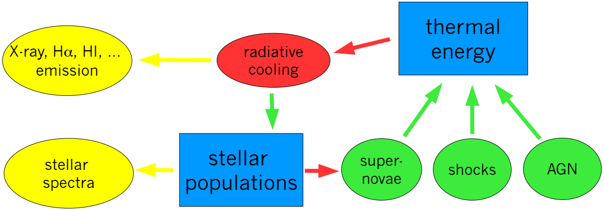
$$\begin{aligned}
\frac{\partial\varepsilon}{\partial t} + \boldsymbol{\nabla}\cdot[(\varepsilon + P_{\mathrm{th}} + P_{\mathrm{cr}})\boldsymbol{v}] &= P_{\mathrm{cr}}\boldsymbol{\nabla}\cdot\boldsymbol{v} - \boldsymbol{v}_{\mathrm{st}}\cdot\boldsymbol{\nabla} P_{\mathrm{cr}} \\
\frac{\partial\varepsilon_{\mathrm{cr}}}{\partial t} + \boldsymbol{\nabla}\cdot[P_{\mathrm{cr}}\boldsymbol{v}_{\mathrm{st}} + \varepsilon_{\mathrm{cr}}(\boldsymbol{v} + \boldsymbol{v}_{\mathrm{st}} + \boldsymbol{v}_{\mathrm{di}})] &= -P_{\mathrm{cr}}\boldsymbol{\nabla}\cdot\boldsymbol{v} + \boldsymbol{v}_{\mathrm{st}}\cdot\boldsymbol{\nabla} P_{\mathrm{cr}} \\
\Longleftrightarrow \quad \frac{\partial\varepsilon_{\mathrm{cr}}}{\partial t} + \boldsymbol{\nabla}\cdot[\varepsilon_{\mathrm{cr}}(\boldsymbol{v} + \boldsymbol{v}_{\mathrm{st}} + \boldsymbol{v}_{\mathrm{di}})] &= -P_{\mathrm{cr}}\boldsymbol{\nabla}\cdot(\boldsymbol{v} + \boldsymbol{v}_{\mathrm{st}})
\end{aligned}$$

Cosmological moving-mesh code AREPO (Springel 2010)



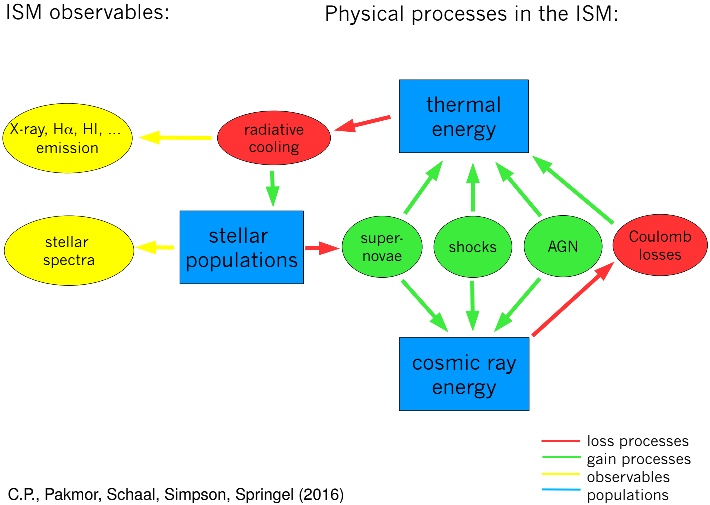
Simulations – flowchart

ISM observables: Physical processes in the ISM:



C.P., Pakmor, Schaal, Simpson, Springel (2016)

Simulations with cosmic ray physics



C.P., Pakmor, Schaal, Simpson, Springel (2016)

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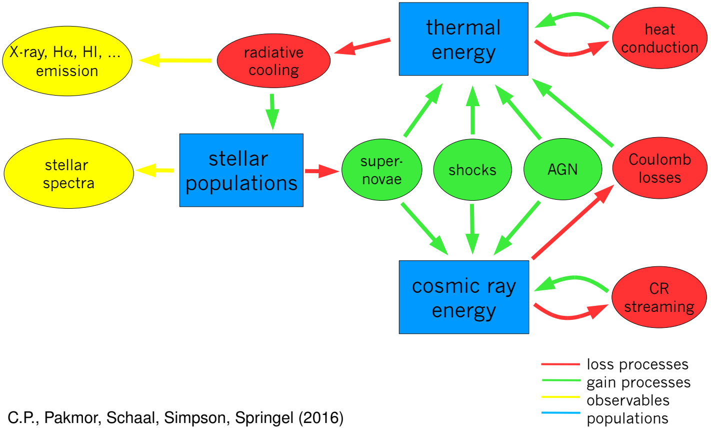

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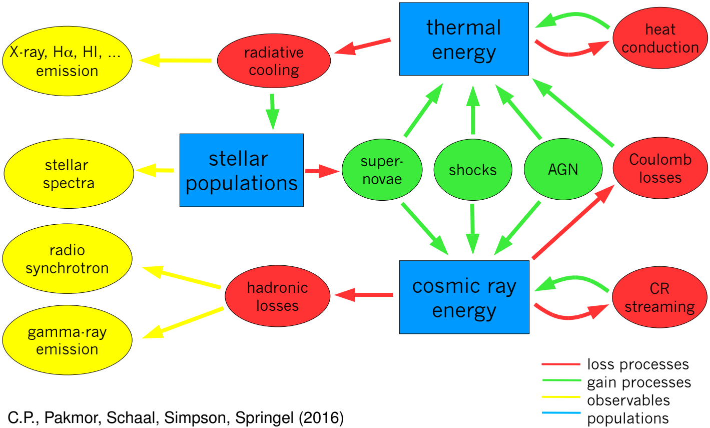
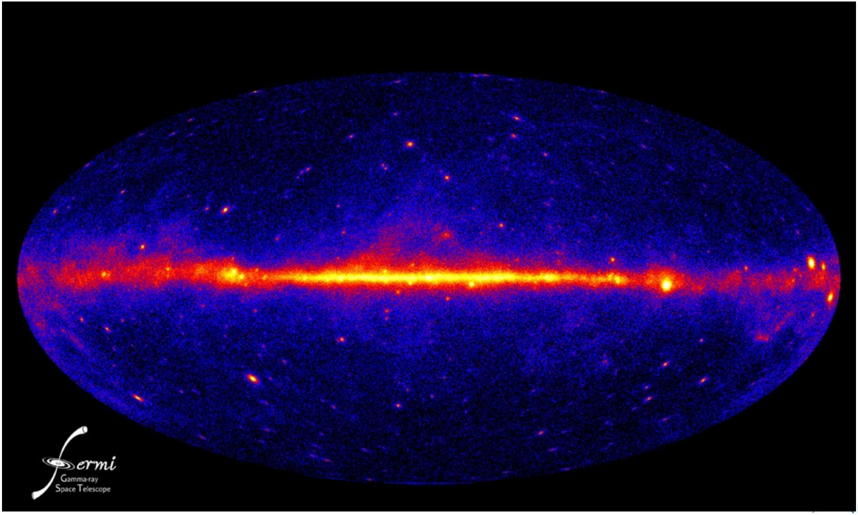


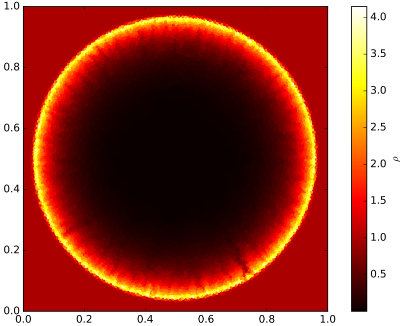
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Gamma-ray emission of the Milky Way

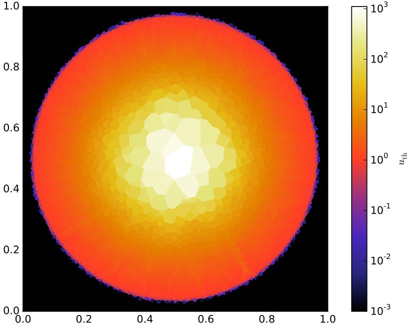


Sedov explosion

density



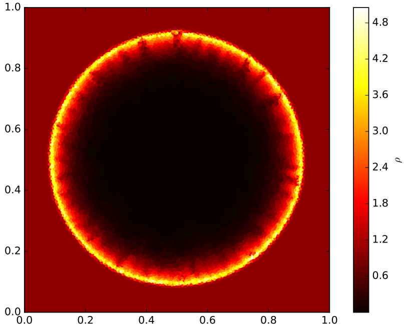
specific thermal energy



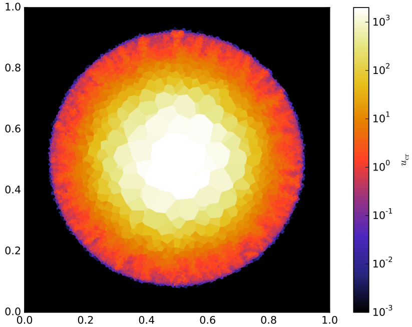
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Sedov explosion with CR acceleration

density



specific cosmic ray energy

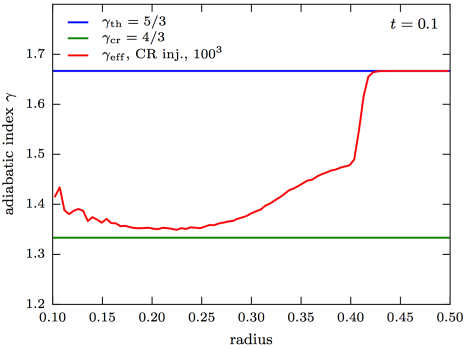


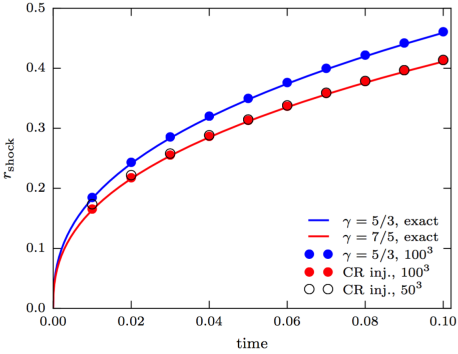
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Sedov explosion with CR acceleration

adiabatic index

shock evolution

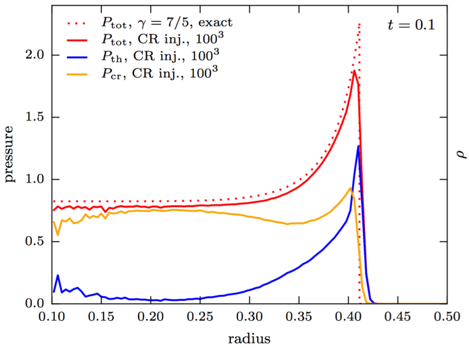




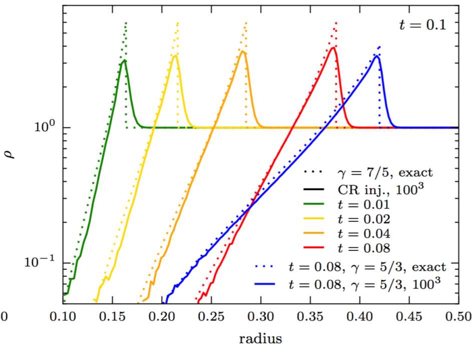
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Sedov explosion with CR acceleration

pressure

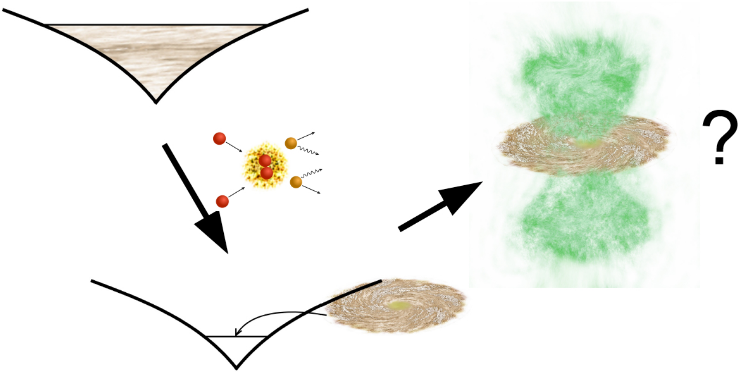


density



C.P., Pakmor, Schaal, Simpson, Springel (2016)

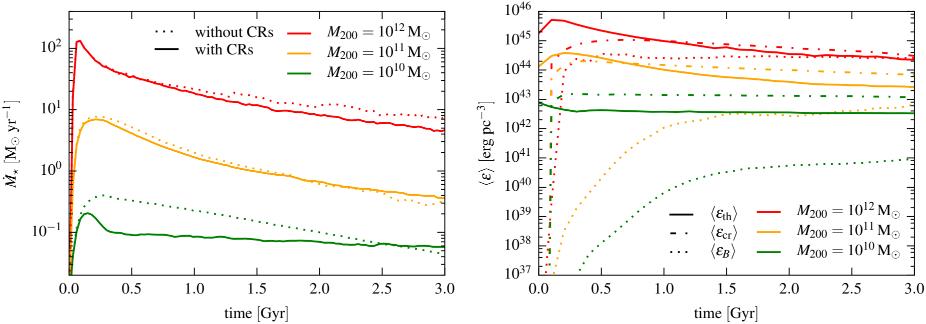
Galaxy simulation setup: 1. cosmic ray advection



C.P., Pakmor, Schaal, Simpson, Springel (2016)
Simulating cosmic ray physics on a moving mesh

MHD + cosmic ray advection

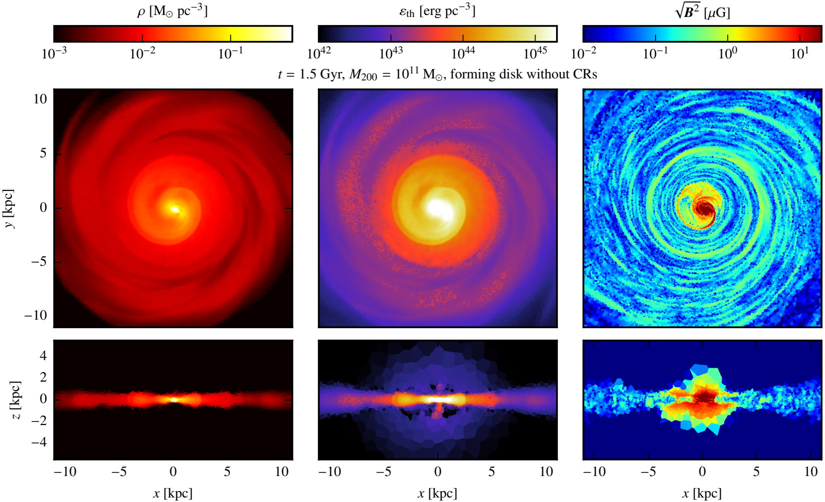
Time evolution of SFR and energy densities



C.P., Pakmor, Schaal, Simpson, Springel (2016)

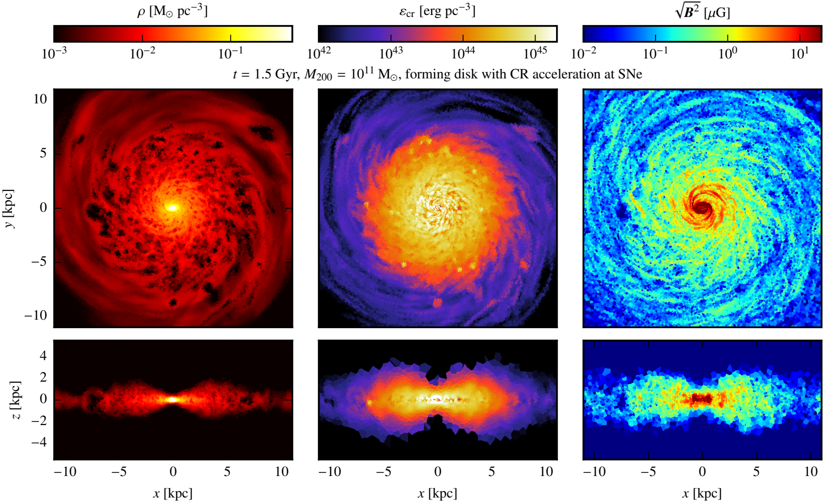
- CR pressure feedback suppresses SFR more in smaller galaxies
- energy budget in disks is dominated by CR pressure
- magnetic dynamo faster in Milky Way galaxies than in dwarfs

MHD galaxy simulation without CRs



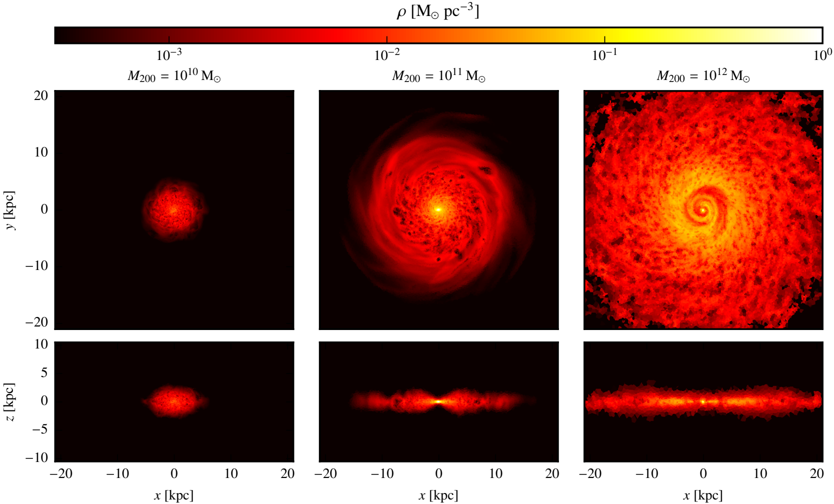
C.P., Pakmor, Schaal, Simpson, Springel (2016)

MHD galaxy simulation with CRs



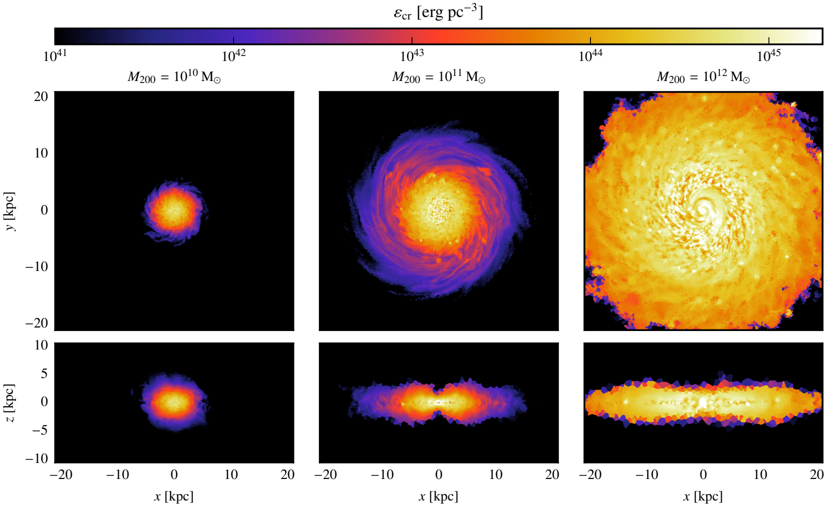
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Gas density in galaxies from 10^{10} to $10^{12}\,\mathrm{M}_\odot$



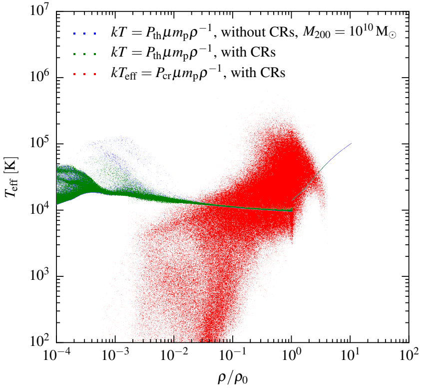
C.P., Pakmor, Schaal, Simpson, Springel (2016)

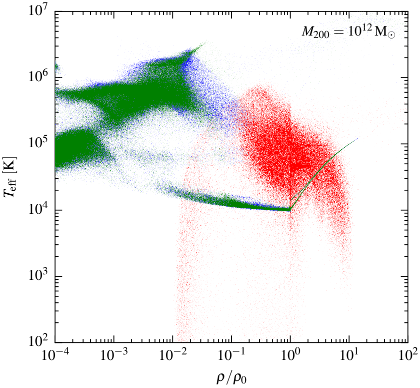
CR energy density in galaxies from 10^{10} to $10^{12}\,\mathrm{M}_\odot$



C.P., Pakmor, Schaal, Simpson, Springel (2016)

Temperature-density plane: CR pressure feedback

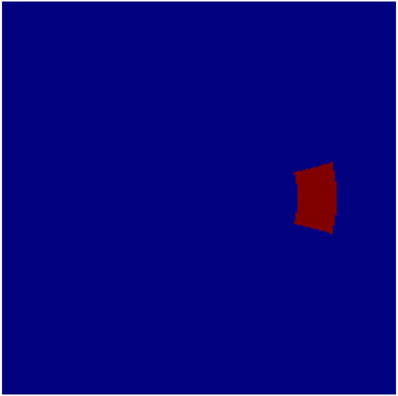




C.P., Pakmor, Schaal, Simpson, Springel (2016)

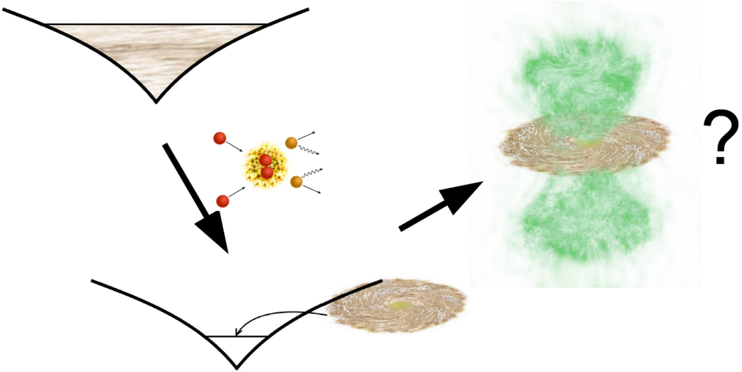
Anisotropic CR diffusion

- diffusion of CR energy density along magnetic field lines
- implemented on unstructured mesh in AREPO
- implicit solver with local time stepping
- obeys 1. and 2. law of thermodynamics (energy and entropy flux conserving)



Pakmor, C.P., Simpson, Kannan, Springel (2016)

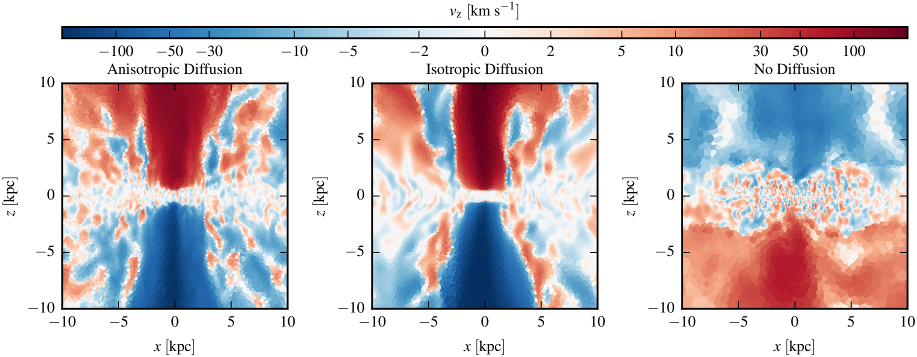
Galaxy simulation setup: 2. cosmic ray diffusion



Pakmor, C.P., Simpson, Springel (2016)
Galactic winds driven by isotropic and anisotropic cosmic ray diffusion in isolated disk galaxies

MHD + cosmic ray advection + diffusion

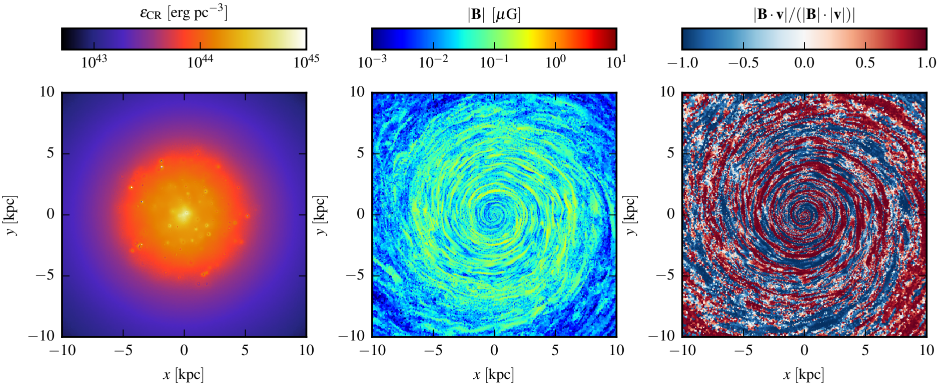
MHD galaxy simulation with CR diffusion



Pakmor, C.P., Simpson, Springel (2016)

- CR diffusion launches powerful winds
- simulation without CR diffusion exhibits only weak fountain flows

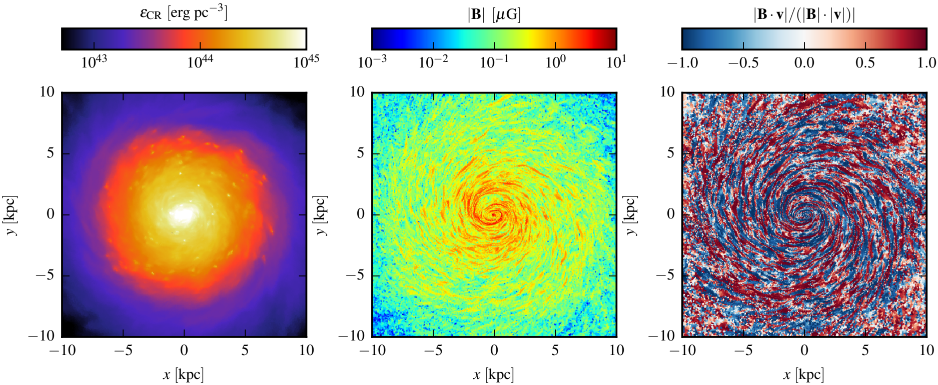
MHD galaxy simulation with CR isotropic diffusion



Pakmor, C.P., Simpson, Springel (2016)

- CR diffusion strongly suppresses SFR
- strong outflow quenches magnetic dynamo to yield $B \sim 0.1\ \mu$G

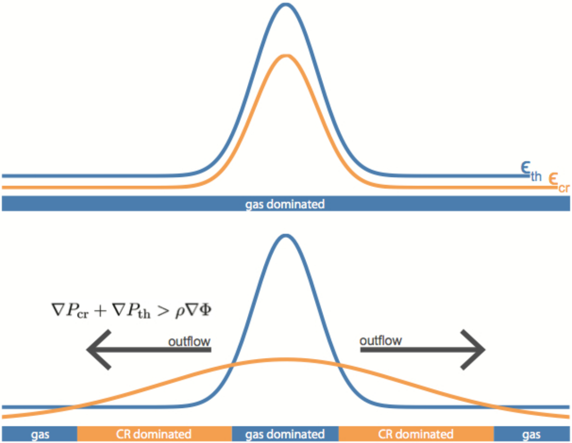
MHD galaxy simulation with CR anisotropic diffusion



Pakmor, C.P., Simpson, Springel (2016)

- anisotropic CR diffusion also suppresses SFR
- reactivation of magnetic dynamo: growth to observed strengths

Cosmic ray driven wind: mechanism



CR streaming: Uhlig, C.P.+ (2012)

CR diffusion: Booth+ (2013), Hanasz+ (2013), Salem & Bryan (2014)

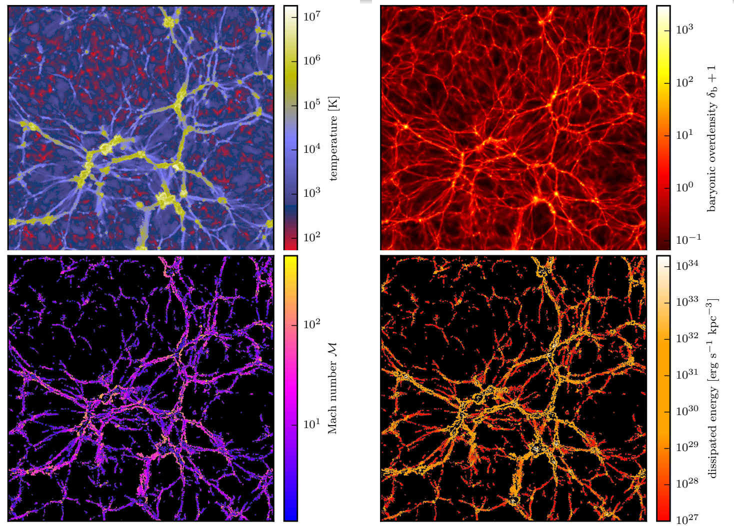
Conclusions on cosmic-ray feedback in galaxies

- CR pressure feedback slows down star formation and provides additional stability to galactic disks
- galactic winds are naturally explained by CR diffusion
- anisotropic CR diffusion necessary for efficient galactic dynamo: observed field strengths of $B \sim 10\,\mu$G

→ versatile CR-MHD code to explore the physics of galaxy formation!

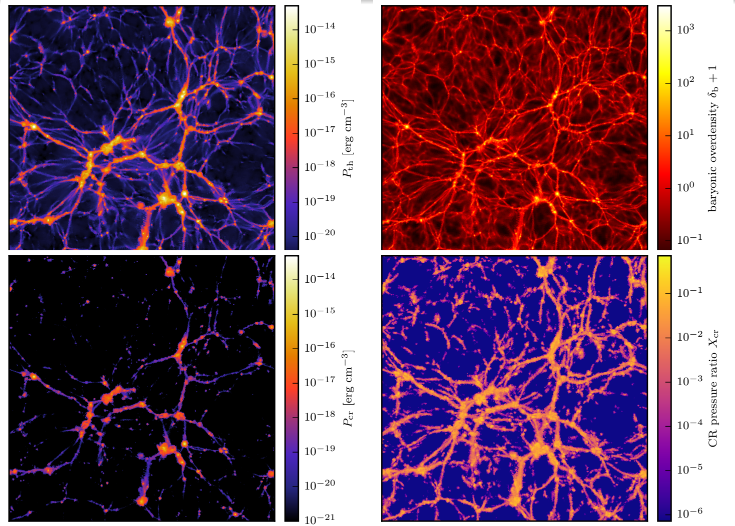
outlook: improved modeling of plasma physics, follow CR spectra, cosmological settings

Cosmological simulations with cosmic rays



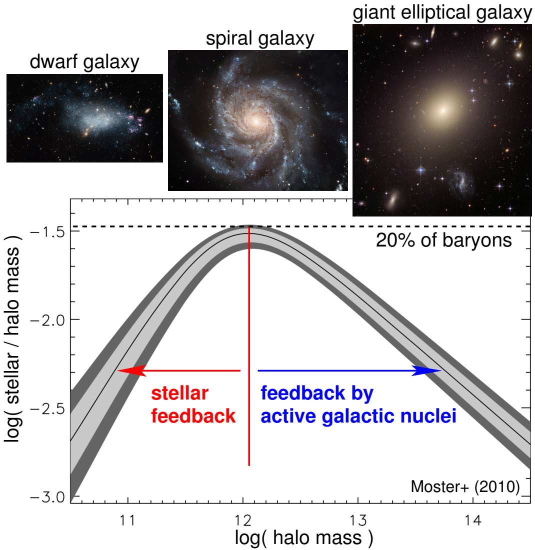
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Cosmological simulations with cosmic rays



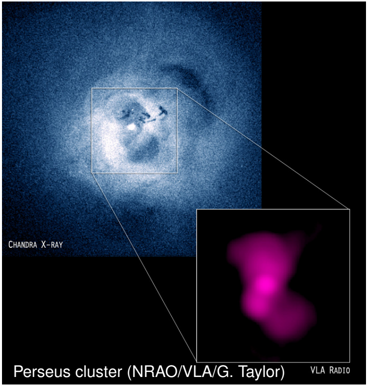
C.P., Pakmor, Schaal, Simpson, Springel (2016)

Puzzles in galaxy formation



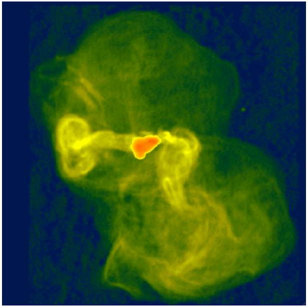
Radio mode feedback by AGN: open questions

- energy source:
  release of non-gravitational accretion energy of a black hole
- jet-ICM interaction and rising bubbles:
  1.) magnetic draping → amplification
  2.) CR confinement vs. release
  3.) excitation of turbulence
- heating mechanism:
  1.) self-regulated to avoid overcooling
  2.) thermally stable to explain T floor
  3.) low energy coupling efficiency
- cosmic ray heating:
  1.) are CRs efficiently mixed into the ICM?
  2.) is the CR heating rate sufficient to balance cooling?
  3.) how universal is this heating mechanism in cool cores?

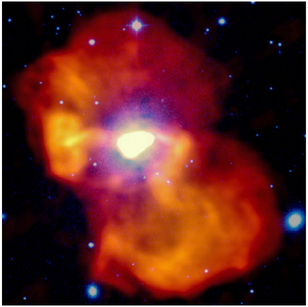


Perseus cluster (NRAO/VLA/G. Taylor)

Messier 87 at radio wavelengths



$\nu = 1.4$ GHz (Owen+ 2000)



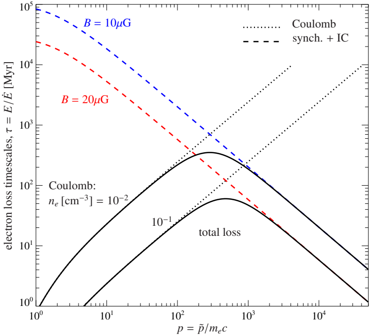
$\nu = 140$ MHz (LOFAR/de Gasperin+ 2012)

- high-ν: freshly accelerated CR electrons
  low-ν: fossil CR electrons → time-integrated AGN feedback!
- LOFAR: halo confined to same region at all frequencies and no low-ν spectral steepening → puzzle of “missing fossil electrons”

Solutions to the “missing fossil electrons” problem

solutions:

- special time: M87 turned on $\sim$ 40 Myr ago after long silence
  $\Leftrightarrow$ conflicts order unity duty cycle inferred from stat. AGN feedback studies (Birzan+ 2012)
- Coulomb cooling removes fossil electrons
  $\rightarrow$ efficient mixing of CR electrons and protons with dense cluster gas
  $\rightarrow$ predicts γ rays from CRp-p interactions:
  $p + p \rightarrow \pi^0 + \ldots \rightarrow 2\gamma + \ldots$



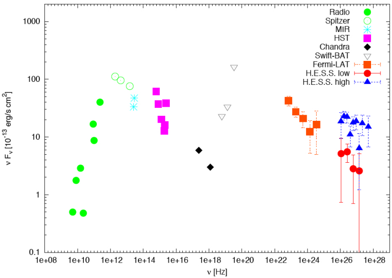
C.P. (2013)

The gamma-ray picture of M87

- high state is time variable
  → jet emission
- low state:
  (1) steady flux

  (2) γ-ray spectral index (2.2)
  $=$ CRp index
  $=$ CRe injection index as probed by LOFAR

  (3) spatial extension is under investigation (?)



Rieger & Aharonian (2012)

→ **confirming this triad would be smoking gun for first γ-ray signal from a galaxy cluster!**

AGN feedback = cosmic ray heating (?)

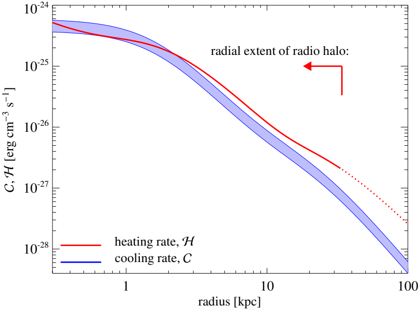
hypothesis: low state γ-ray emission traces π^0 decay within cluster

- cosmic rays excite Alfvén waves that dissipate the energy → heating rate

$$\mathcal{H}_{\mathrm{cr}} = -\boldsymbol{v}_{\mathrm{st}} \cdot \nabla P_{\mathrm{cr}}$$

(Loewenstein+ 1991, Guo & Oh 2008, Enßlin+ 2011, Wiener+ 2013, C.P. 2013)

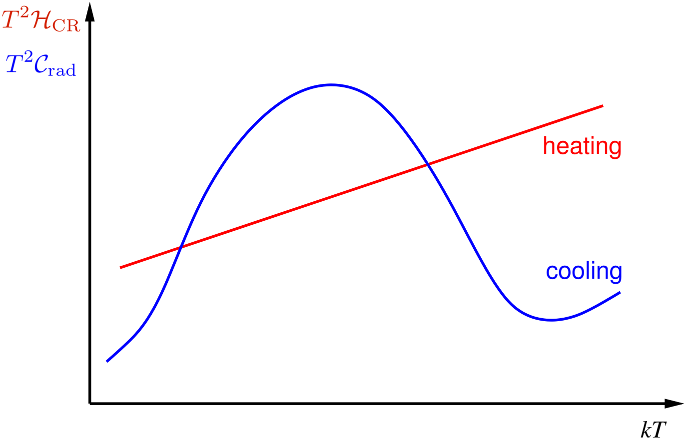
- calibrate P_{cr} to γ-ray emission and $|\boldsymbol{v}_{\mathrm{st}}| = |\boldsymbol{v}_{\mathrm{A}}|$ to radio/X-ray emission
→ spatial heating profile



C.P. (2013)

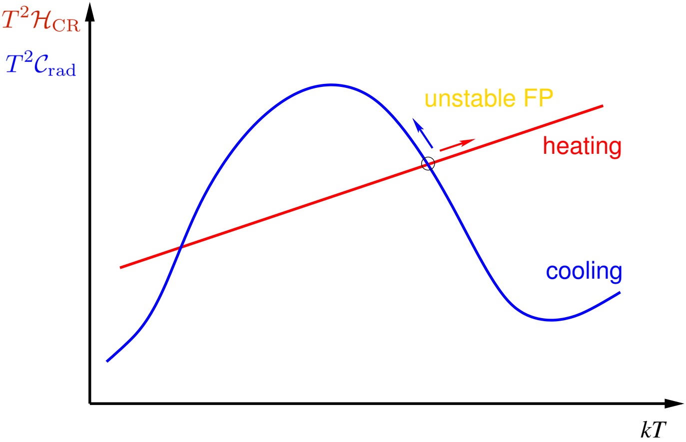
→ cosmic-ray heating matches radiative cooling (observed in X-rays) and may solve the famous "cooling flow problem" in galaxy clusters!

Local stability analysis (1)



- isobaric perturbations to global thermal equilibrium
- CRs are adiabatically trapped by perturbations

Local stability analysis (1)



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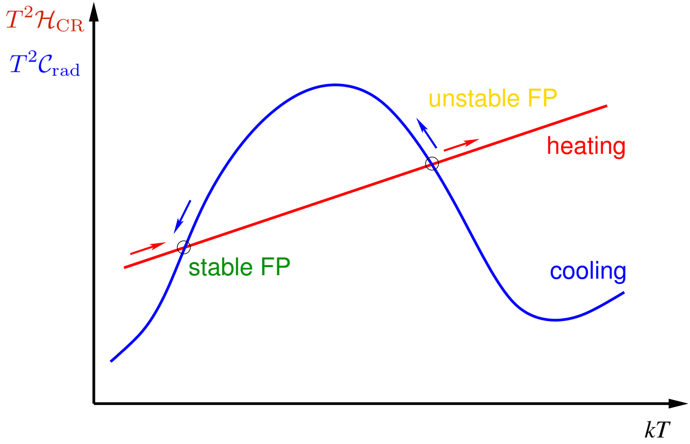

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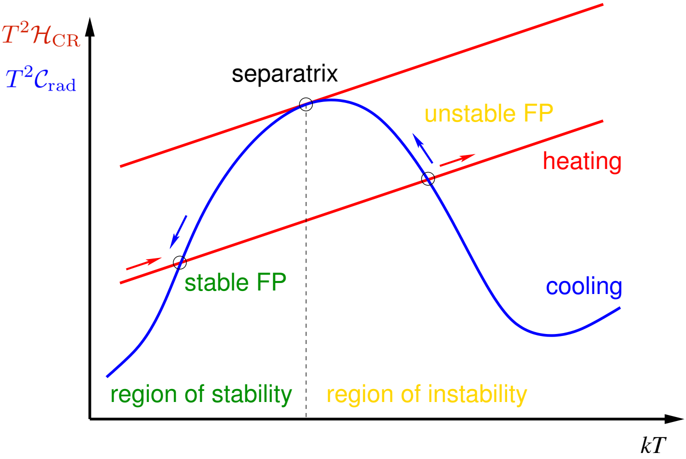
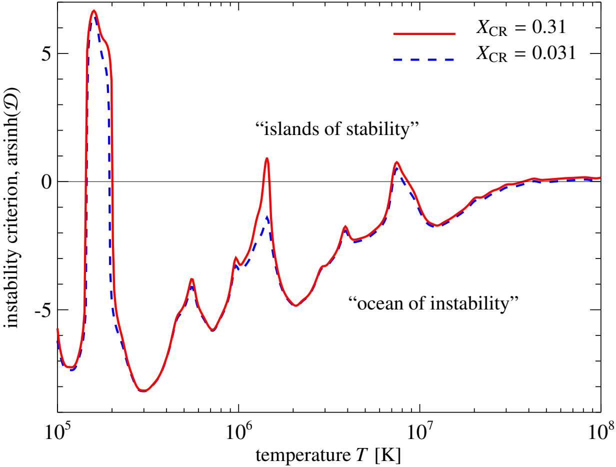


- isobaric perturbations to global thermal equilibrium
- CRs are adiabatically trapped by perturbations

Local stability analysis (2)

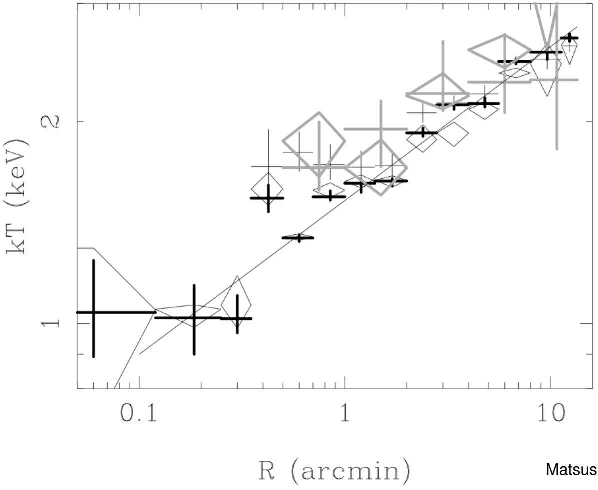
Theory predicts observed temperature floor at $kT \simeq 1$ keV



C.P. (2013)

Virgo cluster cooling flow: temperature profile

X-ray observations confirm temperature floor at $kT \simeq 1$ keV



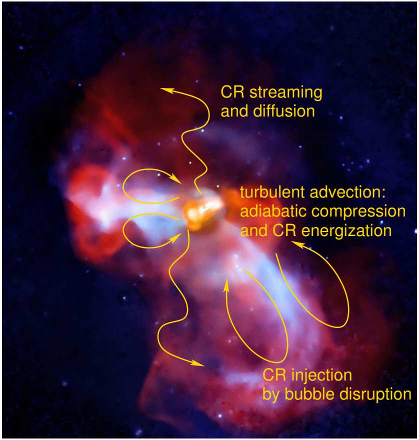
Matsushita+ (2002)

Emerging picture of CR feedback by AGNs

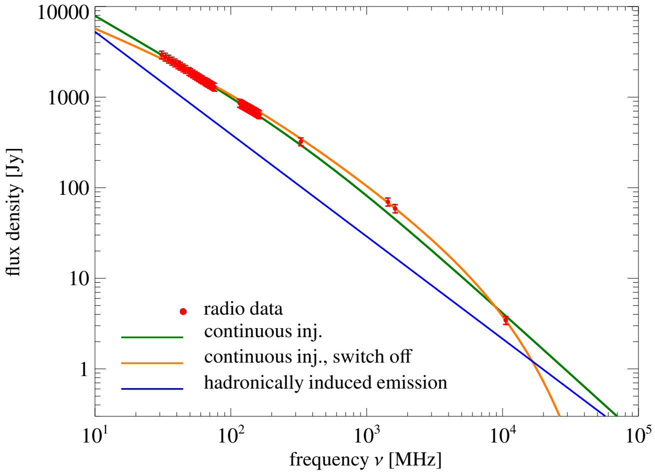
(1) during buoyant rise of bubbles: CRs diffuse and stream outward
→ CR Alfvén-wave heating

(2) if bubbles are disrupted, CRs are injected into the ICM and caught in a turbulent downdraft that is excited by the rising bubbles
→ CR advection with flux-frozen field
→ adiabatic CR compression and energizing: $P_{\rm cr}/P_{\rm cr,0} = \delta^{4/3} \sim 20$ for compression factor $\delta = 10$

(3) CR escape and outward streaming → CR Alfvén-wave heating



Prediction: flattening of high-ν radio spectrum



C.P. (2013)

Conclusions on AGN feedback by cosmic-ray heating

- LOFAR puzzle of "missing fossil electrons" solved by mixing with dense cluster gas and Coulomb cooling
- predicted γ rays identified with low state of M87 $\rightarrow$ estimate CR-to-thermal pressure of $X_{\rm cr} = 0.31$
- CR Alfvén wave heating balances radiative cooling on all scales within the radio halo ($r < 35$ kpc)
- local thermal stability analysis predicts observed temperature floor at $kT \simeq 1$ keV

outlook: couple CRs to AGN jet model, simulate anisotropically steaming CRs, cosmological cluster simulations

need: deeper radio/γ-ray observations + model-informed analysis

CRAGSMAN: The Impact of Cosmic RAys on Galaxy and CluSter ForMAtioN

European Research Council
Established by the European Commission

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Literature for the talk

Cosmic ray feedback in galaxies:

- Pfrommer, Pakmor, Schaal, Simpson, Springel, *Simulating cosmic ray physics on a moving mesh,* 2016, MNRAS.
- Pakmor, Pfrommer, Simpson, Springel, *Galactic winds driven by isotropic and anisotropic cosmic ray diffusion in isolated disk galaxies,* 2016, ApJL.
- Pakmor, Pfrommer, Simpson, Kannan, Springel, *Semi-implicit anisotropic cosmic ray transport on an unstructured moving mesh,* 2016, MNRAS.

AGN feedback by cosmic rays:

- Pfrommer, *Toward a comprehensive model for feedback by active galactic nuclei: new insights from M87 observations by LOFAR, Fermi and H.E.S.S.*, 2013, ApJ, 779, 10.